Thursday, March 16, 2000
POSTER SESSION II
7:00 p.m. UHCL

Mars: The Other Red Meat

Edgett K. S. Malin M. C. Sullivan R. J. Thomas P. Veverka J.
Dynamic Mars: New Dark Slope Streaks Observed on Annual and Decadal Time Scales **[#1058]**
MGS MOC images show new dark mass-wasting deposits formed over periods of 11 Mars years (Viking to MGS) and 1 Mars year (MGS AB-1 to MGS Mapping Phase). Mars is a dynamic planet with geologic processes occurring today.

Figueredo P. H. Greeley R.
Local Variations in Aeolian Deposits in Melas Chasma, Mars **[#1024]**
We studied variations in dune orientation and spacing in a MOC image from Melas Chasma. Using morphometric relationships, we estimated the thickness of the aeolian cover in 337 places within the area and integrated them into an isopach map.

Haldemann A. F. C. Forsberg N. K. Golombek M. P. Bridges N. T.
Far-Field Rock Size-Frequency Distribution at the Mars Pathfinder Landing Site and Comparison to the Near Field **[#1846]**
Detailed measurements of rocks in the far field at the Mars Pathfinder landing site are consistent with the near field exponential drop off in the cumulative number or area covered by large diameter rocks (and with similar behavior at the Viking sites).

Craddock R. A. Golombek M. Howard A. D.
Analyses of Rock Size-frequency Distributions and Morphometry of Modified Hawaiian Lava Flows: Implications for Future Martian Landing Sites **[#1649]**
Both the size-frequency distribution and morphometry of rock populations emplaced by a variety of geologic processes in Hawaii indicate that such information may be useful in planning future landing sites on Mars and interpreting the surface geology.

Stooke P. J.
The Pathfinder Landing Area in MGS/MOC Images **[#1133]**
Features seen in Pathfinder panoramas were looked for in SPO-phase MGS/MOC images. A possible location places Pathfinder 20 m east of the point found by plotting azimuths to horizon features. The method may be useful for locating the Viking landers.

Stoker C. R. Kanefsky B.
The Superresolved Super Pan: Improved Resolution of the Mars Pathfinder Landing Site Using Superresolution on the IMP Super Pan Data Set **[#1753]**
The Imager for Mars Pathfinder Super Pan image cubes were combined using Superresolution to produce a data set with a factor of two higher spatial resolution than the starting products. The data product is presented.

Koehler U. Head J. W. III Hiesinger H. Mustard J. Pratt S. Thomson B.
Geology and Topography in the Sinus Meridiani Region, Proposed as a Possible Mars Landing Site **[#1805]**
A part of the Sinus Meridiani region is under consideration as a possible landing site for a Mars lander mission. Here morphological and topographical observations of this area from Viking, MOC and MOLA data are presented and discussed.

Tanaka K. L. Joyal T. Wenker A.
The Isidis Plains Unit, Mars: Possible Catastrophic Origin, Tectonic Tilting, and Sediment Loading **[#2023]**
We propose that a deposit covering most of Isidis Planitia formed by catastrophic erosion of materials along the western margin of the topographic basin and subsequently deformed by sediment loading and tilting caused by loading of a huge deposit in Utopia Planitia.

Kolb E. J. Tanaka K. L.
Further Volume and Thickness Constraints for the South Polar Layered Deposits of Mars **[#2063]**
Employing Mars Orbital Laser Altimeter data, the volume and thickness of south polar layered deposit material has been further constrained.

Bradley B. A. Grosfils E. B. Sakimoto S. E. H.
Boundaries and Stratigraphy of the Medusae Fossae Formation and Elysium Basin Materials Using Mars Orbiter Laser Altimeter (MOLA) Data **[#2055]**
MOLA data provide a new method for evaluating the accuracy of unit boundaries previously proposed for the Medusae Fossae Formation. Here we examine the MFF deposits along the dichotomy boundary from 202–222W.

Kreslavsky M. A. Head J. W.
Kilometer-scale Roughness of Martian Surface from MOLA Data: Characterization of Geological Units **[#1144]**
The median absolute value of differential slope is used as a measure of surface roughness at 0.6–20 km baselines. Geological units demonstrate distinctive dependences of roughness on the baseline length.