

**THE IMPACT OF SUBSURFACE SOURCES AND SINKS ON THE EVOLUTION OF EARLY ATMOSPHERIC CO<sub>2</sub>, CH<sub>4</sub>, AND D/H.** E. Chassefière<sup>1,2</sup> and F. Leblanc<sup>3</sup>, <sup>1</sup>Univ Paris-Sud, Laboratoire IDES, UMR8148, Université Paris-Sud, Bât. 504, Orsay, F-91405, France, e-mail address : eric.chassefiere@u-psud.fr , <sup>2</sup>CNRS, Orsay, F-91405, France, <sup>3</sup>Laboratoire ATmosphères Milieux Observations Spatiales/IPSL, CNRS-UVSQ, Université Pierre et Marie Curie, Boite 102, 4 Place Jussieu, 75005, Paris.

**Introduction:** Like Venus and Earth, Mars has been endowed with large amounts of CO<sub>2</sub> during main accretion. Thermal escape of carbon is expected to have removed most of the primordial CO<sub>2</sub>, with a characteristic time to lose 1 bar of CO<sub>2</sub> in the range from 1 to 10 Myr : if most of Mars' CO<sub>2</sub> was emplaced during the main accretion phase, the entire inventory could have been lost within 40 Myr [1]. Following this loss to space of CO<sub>2</sub>, a secondary atmosphere progressively formed through volcanic outgassing of CO<sub>2</sub> [2]. According to the morphological analysis, 0.3 bar of CO<sub>2</sub> would have been released during the last 4 Gyr [3]. Crust formation modelling suggests that a total ≈1 bar of CO<sub>2</sub> have been outgassed, mostly during the Noachian [2]. The present CO<sub>2</sub> pressure is 7 mbar, showing that most of the atmospheric CO<sub>2</sub> present at the late Noachian has been removed during subsequent periods, either by escape to space, or by trapping in the subsurface, or both.

Mars has been similarly endowed with large amounts of water during accretion, equivalent to the content of several Earth oceans, corresponding to a several 10 km thick Global Equivalent Layer –GEL– [4]. The present total inventory of water on Mars has been estimated to correspond to a 35 m thick GEL [5]. The mega-regolith capacity is large, with up to ~500 m potentially trapped in the cryosphere, and hypothetically several additional hundreds of meters (up to ~500 m) of ground water surviving at depth below the cryosphere [6]. A ~500 m thick GEL is generally assumed to be required to explain the formation of outflow channels [7]. The fate of this water is unknown. It may have escaped to space, and/or be trapped in the subsurface under the form of hydrated minerals, water ice or, at large depth, liquid water.

The fate of CO<sub>2</sub> (several hundred mbar up to several bar) and H<sub>2</sub>O (several 100 m thick GEL) present in the atmosphere and at the surface (or in the close subsurface) at the late Noachian is poorly constrained, and remains mysterious. Non-thermal escape of carbon and oxygen, driving the losses of, resp., CO<sub>2</sub> and H<sub>2</sub>O, is not efficient enough to have removed hundred millibars of CO<sub>2</sub> and a several hundred meters thick GEL of H<sub>2</sub>O [8,9]. Trapping of these volatiles in the crust under the form, resp., of carbonates and hydrates could have been the major process at the origin of Mars desertification.

**A long-term carbon cycle involving CO<sub>2</sub>, CH<sub>4</sub> and carbonates:** The recent discovery of methane in the Martian atmosphere at a typical 10 ppb level [10,11] suggests that CH<sub>4</sub>, rapidly oxidized to CO<sub>2</sub> in the atmosphere (or the regolith), could have been an important contributor to the atmospheric carbon inventory, in addition to volcanic CO<sub>2</sub> [3]. If CH<sub>4</sub> is the result of serpentinization in crustal hydrothermal systems [12], the carbon in the released methane could originate from subsurface carbonates that were decomposed by hydrothermal fluids. Alternatively, CH<sub>4</sub> may originate in magmatic CO<sub>2</sub> converted to CH<sub>4</sub> through fluid-rock interaction in deep hydrothermal fluids [13]. Assuming that the hydrothermal activity has remained proportional to the extrusion rate of volcanic lava, estimated from existing geomorphological analysis of the Martian surface [3], an upper limit on the cumulated amount (from the late Noachian to the present time) of CO<sub>2</sub> resulting from CH<sub>4</sub> release in the range from 0.2-2 bar, of the same order as the volcanic CO<sub>2</sub>, has been proposed [8]. Because non-thermal escape cannot have removed more than ≈10 mbar of CO<sub>2</sub> since the late Noachian [8], it may be thought that atmospheric carbon has been significantly cycled to the crust at late Noachian, Hesperian and Early Amazonian through subsurface hydrological activity. In this way, the production of CO<sub>2</sub> through CH<sub>4</sub> release and further oxidation, and the removal of CO<sub>2</sub> from the atmosphere, could have a common origin and be two facets of a currently, although progressively damping with time, active hydrological system.

Such a long term carbon cycle, with a progressive net removal of CO<sub>2</sub> from the atmosphere and subsequent carbonate deposition in the subsurface, would explain the present low value of the amount of CO<sub>2</sub> contained in the atmosphere (7 mbar) and polar reservoirs (4-5 mbar [14]), compared to the ≈1 bar CO<sub>2</sub> secondary atmosphere produced by volcanic outgassing. Because a C atom could have been cycled several times through the crust since its release from the mantle by volcanism, the cumulated CH<sub>4</sub> released rate (up to 0.2-2 bar) must not be interpreted as the content of an isolated subsurface reservoir. It rather suggests that an efficient carbon cycle has been maintained by hydrothermal processes, with a substantial fraction of the volcanic outgassed carbon being cycled one or several times through crustal carbonates. Contrary to the steady state carbon cycle at work on Earth, a progressive damping of the carbon cycle has occurred on Mars due to the

