

**PETROGRAPHY AND GEOCHEMISTRY OF METALS IN ALMAHATA SITTA UREILITES.**

A. J. Ross<sup>1,2</sup>, J. S. Herrin<sup>3</sup>, D. W. Mittlefehldt<sup>3</sup>, H. Downes<sup>1,2</sup>, C. L. Smith<sup>1,4</sup>, M. R. Lee<sup>4</sup>, A. P. Jones<sup>2</sup>, P. Jenniskens<sup>5</sup> and M. H. Shaddad<sup>6</sup>. <sup>1</sup>IARC, Department of Mineralogy, The Natural History Museum, London, UK ([aidan.ross@ucl.ac.uk](mailto:aidan.ross@ucl.ac.uk)), <sup>2</sup>Centre for Planetary Sciences, Joint UCL/Birkbeck Research School of Earth Sciences, London, UK, <sup>3</sup>NASA Johnson Space Center, Houston, TX 77058, USA, <sup>4</sup>School of Geographical and Earth Sciences, University of Glasgow, Glasgow, UK, <sup>5</sup>SETI Institute, Mountain View, CA 94043, USA, <sup>6</sup>Physics Dept., University of Khartoum, Khartoum 11115, Sudan.

**Introduction:** Ureilites are ultramafic achondrites, predominantly composed of olivine and pyroxenes with accessory carbon, metal and sulfide. The majority of ureilites are believed to represent the mantle of the ureilite parent body (UPB) [1]. Although ureilites have lost much of their original metal [2], the metal that remains retains a record of the formative processes. Almahata Sitta is predominantly composed of unbrecciated ureilites with a wide range of silicate compositions [3,4]. As a fall it presents a rare opportunity to examine fresh ureilite metal in-situ, and analyzing their highly siderophile element (HSE) ratios gives clues to their formation. Bulk siderophile element analyses of Almahata Sitta fall within the range observed in other ureilites [5]. We have examined the metals in seven ureilitic samples of Almahata Sitta (AS) and one associated chondrite fragment (AS#25).

**Methods:** *Electron-beam techniques:* Imaging (using a LEO 1455VP SEM with Oxford Instruments INCA software) and major element composition analyses (using a WDS Cameca SX100 EMPA) were conducted at the NHM. EBSD analysis was conducted at the University of Glasgow using an EDAX-TSL system attached to a Quanta 200F field-emission SEM that was operated in low vacuum mode.

*LA-ICP-MS:* Analyses to obtain trace element compositions were performed at NASA JSC using a New Wave UP-193 SS laser ablation system attached to a Thermo Element2-XR ICP-MS. Laser power was 1-2 GW/cm<sup>2</sup>, with a repetition rate of 10 Hz. Spot sizes of 20 to 100 µm were used, depending on metal grain size. Some grains were too small to analyze. The Filomena and Hoba iron meteorites, NIST 610 glass, and NBS 1168 and NBS 1178 steels were used as standards. Ni values from EPMA were used for internal calibration for all samples except AS#15, #25 and #33 where Fe+Co+Ni were summed to 100 wt%. The following isotopes were analyzed: <sup>29</sup>Si, <sup>31</sup>P, <sup>34</sup>S, <sup>53</sup>Cr, <sup>54</sup>Fe, <sup>55</sup>Mn, <sup>57</sup>Fe, <sup>59</sup>Co, <sup>60</sup>Ni, <sup>61</sup>Ni, <sup>62</sup>Ni, <sup>63</sup>Cu, <sup>65</sup>Cu, <sup>66</sup>Zn, <sup>69</sup>Ga, <sup>72</sup>Ge, <sup>74</sup>Ge, <sup>75</sup>As, <sup>95</sup>Mo, <sup>101</sup>Ru, <sup>103</sup>Rh, <sup>105</sup>Pd, <sup>106</sup>Pd, <sup>121</sup>Sb, <sup>182</sup>W, <sup>185</sup>Re, <sup>192</sup>Os, <sup>193</sup>Ir, <sup>194</sup>Pt, <sup>195</sup>Pt and <sup>197</sup>Au.

**Petrography and chemistry:** Metals in AS occur as interstitial grain-boundary vein metal, inclusions in silicates, associated with carbon, or in reduction rims, as seen in other unbrecciated ureilites [6]. Reduction rim metal in silicates is not discussed further here as it is a secondary product [6,7]. Some AS samples contain significantly more metal than seen in “typical”

ureilites. Grain boundary metal, whilst most common, is very fine-grained and as such causes significant problems for spot HSE analysis. Sulfide (as troilite) is present in all samples. High-Si metals [8] (e.g. suessite) have not been found in AS. Metal compositions are reported in weight % unless otherwise stated.

*AS#7:* is an anomalous fragment, described as porous [4] with a high percentage of metal. Included metals occur in both olivine and low-Ca pyroxene (LCP), although interstitial metal grains and those associated with porous areas are much larger. Compositional variation (Ni~1.7-3.8%, Co~0.2-0.3%, Si~<dt-1.9%, P~0.2-1.3%) is smaller than found for other metal-rich samples, but a wide range in Ir/Pd~1.1-8.5 is observed.

*AS#15:* is an augite-bearing ureilite with LCP (Wo~5/En~84) and olivine (Fo~92) [4]. Most metal occurs as interstitial veins, with one yielding Ir/Pd~1.4.

*AS#22:* shows a typical ureilitic texture with olivine (Fo~80.3), and two distinct LCPs (Wo~4.7/En~79.0 and Wo~9.5/En~73.9). Non-reduction metal occurs only along grain boundaries with varying compositions (Ni~2.8-5.9%, Co~0.2-0.4%, Si~0.02-2.2%, P~0.1-0.3%). Ir/Pd varies from 2.2 to 3.2.

*AS#25:* is an H5 chondrite collected within the AS fall area, although the relationship with AS ureilites is unclear [4]. Metal is unweathered. Ir/Pd~0.8-1.1.

*AS#27:* consists of olivine (Fo~85.2) and LCP (Wo~5.1/En~81.5). Some metal has unusually high Ni contents (up to 11.3%) compared to other ureilites (<6% [6]), with grains included in LCP being highest. High-Ni metals have previously been reported in polymict ureilite breccias [8] but those may have an exogenic origin. The range of metal compositions (Ni~5.0-11.3%, Co~0.3-0.5%, Si~0.02-0.3%, P~0.2-1.34%) is widest in the LCP-included metals. Those at grain boundaries and associated with carbon have a smaller range in Ni (5.3-5.8%). HSE ratios are closest to chondritic in this sample, with Ir/Pd~0.9-1.7.

*AS#33:* is made up of multiple small (~1mm) fragments, only some of which have silicate attached. Most of the mass consists of large metal blebs (some >200 µm in diameter) with kamacite cores and taenite rims, in sulfide. Fine-scaled metal dendrites are present in some areas of the sulfide. This sample is similar to metal-rich areas in AS sample MS-166 reported in [3].

