

### REFLECTANCE SPECTRA COMPARISONS OF AN ALLENDE DARK INCLUSION AND ASTEROIDS

P. C. Buchanan<sup>1</sup>, M. E. Zolensky<sup>2</sup>, T. Hiroi<sup>3</sup>, L. Le<sup>4</sup>, and C. Galindo Jr.<sup>5</sup>. <sup>1</sup>Kilgore College, Kilgore, TX 75662. E-mail: pbuchanan@kilgore.edu. <sup>2</sup>Astromaterials Research and Exploration Science KT, NASA Johnson Space Center Houston, TX 77058. <sup>3</sup>Department of Geological Sciences, Brown University, Providence, RI 02912. <sup>4</sup>Jacobs Technology, JSC-KR, Houston, TX 77058. <sup>5</sup>MEI Technologies, NASA-JSC, Houston, TX 77058.

**Introduction:** Several studies have previously described the petrography, mineral chemistry, and petrology of dark inclusions from Allende [e.g., 1,2,3,4]. Increasing numbers of high-quality reflectance spectra of K-complex asteroids, which have been suggested to represent various types of carbonaceous chondrite meteorites, have recently become available [e.g., 5]. Hence, it seemed important to us to compare these data with laboratory-derived reflectance spectra and mineralogical data of Allende dark inclusions to determine if any of these clasts might serve as petrologic analogues for asteroids.

**Data:** MZA is a large Type B dark inclusion composed of chondrule-like aggregates of fine-grained olivine in an opaque matrix that also is in large part composed of fine-grained olivine with minor pyroxene and opaques. Type B dark inclusions have been interpreted as the product of aqueous alteration of CV-type material followed by dehydration [e.g., 2,3,4]. Other workers [e.g., 1] have interpreted these clasts as the result of nebular processes. Electron microprobe analyses of fine-grained olivine from clast MZA indicate an average olivine composition of Fo<sub>58.2</sub>. This agrees well with X-ray diffraction analysis of a bulk sample of the dark inclusion, which indicates an approximate average olivine composition of Fo<sub>54</sub>.

The laboratory-derived reflectance spectrum of MZA has a distinctive M2 band at approximately 1.05 $\mu$ m. Based on the work of Sunshine and Pieters [5], the location of this stable M2 band is consistent with an average olivine composition of Fo<sub>50-60</sub>. This laboratory-derived spectrum is distinctly different from that of Allende and some other CV3 meteorites [e.g., 6]. Among the K-complex asteroids, this spectrum is most similar to that of 661 Cloelia [6]. Ironically, the spectrum of MZA is also very similar to that of the A-type asteroid 446 Aetemitas with almost identical location of M2 band and similar spectral features [7,8].

**References:** [1] Johnson C. A. et al. 1990. *Geochimica et Cosmochimica Acta* 54:819-830. [2] Kojima T. and Tomeoka K. 1996. *Geochimica et Cosmochimica Acta* 30:2651-2666. [3] Krot A. N. et al. 1995. *Meteoritics* 30:748-775. [4] Buchanan P. C. 1997. *Geochimica et Cosmochimica Acta* 61:1733-1743. [5] Sunshine J. M. and Pieters C. M. 1998. *Journal of Geophysical Research* 103:13,675-13,688. [6] Clark B. E. et al. 2009. *Icarus* 202:119-133. [7] Bus S. J. and Binzel R. P. 2002. *Icarus* 158:106-145. [8] Burbine et al. 2008. *Reviews in Mineralogy and Geochemistry* 68:273-343.