

OXYGEN ISOTOPIC RATIOS RESULTING FROM CO SELF-SHIELDING. J. R. Lyons¹, ¹Earth & Space Sciences, UCLA, Los Angeles, CA 90095-1567; jimlyons@ucla.edu.

Introduction: A key issue with the CO self-shielding hypothesis [1] as the explanation for the oxygen isotope distribution of the solar system is the resulting $\delta^{17}\text{O}/\delta^{18}\text{O}$ value of the product oxygen. Simply put, does CO self-shielding yield a $\delta^{17}\text{O}/\delta^{18}\text{O}$ ratio that agrees with the CAI mixing line to within a few percent?

Most self-shielding models [2-4] have assumed equal or nearly equal fractionation in ^{17}O and ^{18}O during CO photodissociation. These papers relied on CO shielding functions [5] as a means of avoiding the complete integration over wavelength of the CO isotopic cross sections, but shielding functions were evaluated for only C^{16}O and C^{18}O in [5]. The shielding functions for CO isotopologues have recently been updated [6], and now include C^{17}O . Molecular constants for C^{17}O in the FUV range (91-108 nm) are unmeasured, so the C^{17}O shielding functions are necessarily approximate.

Measurements of isotope effects during CO photodissociation have been performed at the ALS synchrotron [7], and have yielded a wavelength dependence in the $\delta^{17}\text{O}/\delta^{18}\text{O}$ ratio. At long wavelengths (105 & 107.6 nm), they observed $\delta^{17}\text{O}/\delta^{18}\text{O} \sim 1.4$ -1.8, and at short wavelengths (92 & 94 nm) they saw $\delta^{17}\text{O}/\delta^{18}\text{O} \sim 0.6$. The experiments in [7] were performed with very high CO column density, such that C^{16}O had an optical depth up to 10^4 at 107.6 nm. This is not an astrophysically relevant scenario (because of dust absorption), but it does illustrate the effect of simultaneous C^{16}O and C^{18}O self-shielding at long wavelengths to yield $\delta^{17}\text{O}/\delta^{18}\text{O} > 1$. At short wavelengths, the shorter lifetimes of the CO excited states causes some line broadening and band overlap, with a resulting mass-dependent fractionation superimposed upon the mass-independent fractionation from discrete self-shielded lines. The resulting slopes will be < 1 as seen by [7].

Astronomical measurements of CO isotopologues in disks [8] and molecular clouds [9] demonstrate that CO self-shielding does occur in these environments, but the precision of the measurements is too low to be a test of CO self-shielding as an explanation for the meteorite oxygen isotope line.

Discussion: Here I will focus on two other factors that affect the $\delta^{17}\text{O}/\delta^{18}\text{O}$ ratio in CO self-shielding, namely excitation temperature and absorption by dust. Figure 1 shows solar nebula model results at CO excitation temperatures of 50 K and 5 K (the assumed H_2 excitation temperatures are 354 K and 11 K). The cal-

culations use C^xO shielding functions from [6], and assume a viscosity $\alpha = 10^{-2}$. The 50 K results for dust uniformly distributed in the gas yield $\delta^{17}\text{O}/\delta^{18}\text{O} \sim .88$, too low to explain the CAI line. Allowing dust to settle yields a slight increase in $\delta^{17}\text{O}/\delta^{18}\text{O}$ due to C^{18}O self-shielding. Shielding functions at 5 K yield $\delta^{17}\text{O}/\delta^{18}\text{O} \sim 1.2$ -1.3, depending on the assumed dust absorption parameter. Clearly, at 5 K CO will be frozen out in the nebula, so higher temperatures are being evaluated. The implication, however, is that self-shielding can yield $\delta^{17}\text{O}/\delta^{18}\text{O}$ values consistent with CAIs.

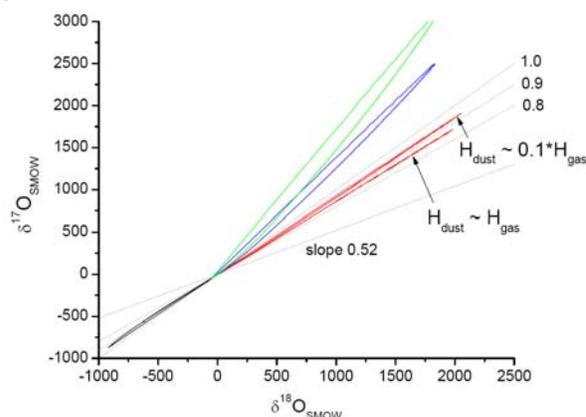


Fig. 1. Nebula model calculations at 30 AU at the midplane for $\text{H}_2\text{O}_{\text{ice}}$. The curves in red use 50 K CO shielding functions (Doppler width 0.3 km s^{-1}) from [6]. Dust is 0.1 microns and is distributed with the gas ($H_{\text{dust}} \sim H_{\text{gas}}$) or is assumed to be settled ($H_{\text{dust}} \sim 0.1 * H_{\text{gas}}$). The curves in green and blue use 5 K shielding functions from [6] (again $b = 0.3 \text{ km s}^{-1}$), and for a dust absorption parameter $\gamma = 1.0$ (green) and 3.5 (blue). At 5 K CO will be frozen out, and so will not undergo gas-phase self-shielding, but the figure illustrates that a variety of $\delta^{17}\text{O}/\delta^{18}\text{O}$ ratios are possible from CO self-shielding. Other temperatures are being evaluated.

References: [1] Clayton R. N (2002) *Nature*, 415, 860-861. [2] Yurimoto H. & Kuramoto K. (2004) *Science*, 305, 1763-1766. [3] Lyons J. R. and Young E. D. (2005) *Nature*, 435, 317-320. [4] Lee J.-E. (2008) *Meteor. Planet. Sci.*, 43, 1351-1362. [5] van Dishoeck E. & Black J. (1988), *Astrophys. J.*, 334, 771. [6] Visser et al. (2009), *Astron & Astrophys.*, 503, 323. [7] Chakraborty S. et al. (2008) *Science* 321, 1328-1331. [8] Smith R. et al. (2009) *Astrophys. J.*, 701, 163-175. [9] Sheffer et al. (2002), *Astrophys. J.*, 574, L171.