

Scientific Goals, Objectives, and Investigations for Venus Exploration (2019)

Draft for Community Feedback (May 2019)

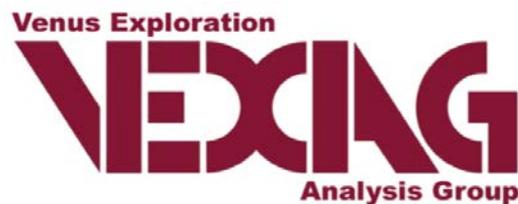


Table of Contents

1. Executive Summary	1
2. VEXAG Goals, Objectives, and Investigations (GOI)	2
3. Descriptions of the Goals, Objectives, and Investigations	4
3.1. Goal I: Understand Venus' early evolution and potential habitability to constrain the evolution of Venus-size (exo)planets	4
3.2. Goal II: Understand atmospheric dynamics and composition on Venus	8
3.3. Goal III: Understand the geologic history preserved on the surface of Venus	12
4. Conclusions	15
Appendix 1: Future Investigations	16
Appendix 2: Linking the 2019 and 2016 VEXAG GOI Documents	17
Appendix 3: Linking the 2019 VEXAG GOI to the Roadmap	19
Appendix 4: List of Acronyms	21

At the VEXAG meeting in November 2017, it was resolved to update the scientific priorities and strategies for Venus exploration. Here we present the 2019 VEXAG Goals, Objectives and Investigations for Venus Exploration.

Prepared by the VEXAG GOI Focus Group: Joseph O'Rourke and Allan Treiman (Co-Chairs), Giada Arney, Paul Byrne, Lynn Carter, Darby Dyar, James Head III, Candace Gray, Stephen Kane, Walter Kiefer, Kevin McGouldrick, Laurent Montesi, Chris Russell, and Suzanne Smrekar.

1. Executive Summary

Venus and Earth are often described as twins. Their sizes and densities are nearly identical, and they are much larger than other terrestrial planetary bodies. Yet past exploration missions reveal that Venus is hellishly hot, devoid of oceans, apparently lacking plate tectonics, and bathed in a thick, reactive atmosphere. A less Earth-like environment is hard to imagine. When and why did Venus and Earth’s evolutionary paths diverge? Did Venus ever host habitable conditions? These fundamental and unresolved questions drive the need for vigorous new exploration of Venus. The answers are central to understanding Venus in the context of terrestrial planets and their evolutionary processes. Critically, Venus provides important clues to understanding our planet—does hot, dry Venus represent the once and future Earth? Current and future efforts to identify planetary systems beyond our Solar System (e.g., the Kepler mission and the Transiting Exoplanet Survey Satellite) are aimed at finding Earth-size planets in the “habitable zone” of their parent stars. Precisely because it may have begun so like Earth, yet evolved to be so different, Venus is the planet most likely to yield new insights into the conditions that determine whether a Venus-sized exoplanet can sustain long-lived habitability.

The planetary science community has consistently identified Venus as a high-priority destination for scientific exploration. In the latest decadal survey (*Visions and Voyages for Planetary Science in the Decade 2013–2022*, National Research Council), Venus was listed as an “important object of study” in all three crosscutting themes (building new worlds, planetary habitats, and workings of solar systems). The decadal committee recommended the Venus In Situ Explorer as one of seven candidate New Frontiers Missions 4 and 5 and the Venus Climate Mission as one of five candidate flagship missions. The midterm review of NASA’s progress towards implementing the decadal survey found that programmatic balance in selected missions is vital to achieving comparative planetology investigations recommended by the decadal survey.

Exciting Venus research has continued since the decadal survey. In particular, the VIRTIS instrument on Venus Express provided tantalizing evidence that the tesserae terrains are felsic—suggesting that they formed in the presence of a large amount of liquid water. Laboratory simulations show that sites of plume-induced subduction on Venus could serve as an analog for early Earth when plate tectonics first began. The JAXA Akatsuki mission has revealed fascinating features in the atmosphere such as planetary-scale standing gravity waves at the cloud tops that are associated with specific topographic features and local times. Conditions similar to those that led to the emergence of life on Earth may have occurred on Venus, but the surface is far too hot for terrestrial organisms today and the colder clouds are extremely acidic.

Through an extended process including input from the community at town hall meetings, VEXAG has developed this list of scientific Goals, Objectives, and Investigations to both serve the priorities of the latest decadal survey and motivate future community efforts. In particular, NASA’s future exploration of Venus should strive to achieve three non-prioritized Goals:

- I. Understand Venus’ early evolution and potential habitability to constrain the evolution of Venus-sized (exo)planets,
- II. Understand atmospheric composition and dynamics on Venus, and
- III. Understand the geologic history preserved on the surface of Venus and the present-day couplings between the surface and atmosphere.

This document describes six related Objectives and twenty-three associated Investigations. Every Investigation was judged as technically and programmatically feasible and of very high scientific value. Ultimately, a sustained program of Venus exploration would unveil the workings of Earth’s nearest neighbor with broad scientific implications for our Solar System and beyond.

2. VEXAG Goals, Objectives, and Investigations (GOI)

Table 1 summarizes this entire report. Understanding Venus as a planetary system requires making progress in many scientific areas. Therefore, Goals and Objectives are not prioritized. Investigations are categorized as either Essential (1), Important (2), or Targeted (3) based on their relationship to the overarching Objective. Completing all Essential (1) Investigations essentially fulfills the Objective. Important (2) Investigations address many aspects of the Objective and provide valuable context for other Investigations. Targeted (3) Investigations address particular aspects of an Objective that significantly contribute to our overall understanding of Venus. Crucially, Investigations are not prioritized in any way within each category, meaning that all Essential Investigations may be described as of the highest priority.

Potential Investigations that were judged as having less than high scientific value are omitted entirely from this report. Investigations that would have high merit but are not technically feasible within the timescale of the VEXAG “Roadmap for Venus Exploration” are not included in Table 1, although some possibilities are discussed in Appendix 1 of this report.

Investigations from the 2016 VEXAG GOI are included in the current version. However, the 2016 Goals focused separately on 1) the atmosphere, 2) surface and interior processes, and 3) the atmosphere-surface interface. The 2019 VEXAG GOI blends Investigations of different focus areas to achieve big-picture Goals and has been iterated with other VEXAG focus groups to serve as the foundation for the VEXAG Roadmap and Technology reports.

Goal	Objective	Investigation
I. Understand Venus’ early evolution and potential habitability to constrain the evolution of Venus-size (exo)planets.	A. Did Venus have temperate surface conditions and liquid water at early times?	HO. Hydrous Origins (1). Determine whether Venus shows evidence for abundant silicic igneous rocks and/or ancient sedimentary rocks.
		RE. Recycling (1). Search for structural, geomorphic, and chemical evidence of crustal recycling on Venus.
		AL. Atmospheric Losses (2). Quantify the processes by which the atmosphere of Venus loses mass to space, including interactions between magnetic fields and incident ions and electrons.
		MA. Magnetism (3). Characterize the distribution of any remanent magnetism in the crust of Venus.
	B. How does Venus elucidate possible pathways for planetary evolution in general?	IS. Isotopes (1). Measure the isotopic ratios and abundances of D/H, noble gases, oxygen, nitrogen, and other elements in the atmosphere of Venus.
		LI. Lithosphere (1). Determine lithospheric parameters on Venus that are critical to rheology and potential geodynamic transitions, including: stress state, water content, physical structure, and elastic and mechanical thicknesses.
		HF. Heat flow (2). Determine the thermal structure of the lithosphere of Venus at present day and measure in situ heat flow.
		CO. Core (2). Measure the size of the core of Venus and determine whether it remains partially liquid.

Table 1 (continued). VEXAG Goals, Objectives, and Investigations		
II. Understand atmospheric dynamics and composition on Venus.	A. What processes drive the global atmospheric dynamics of Venus?	DD. Deep Dynamics (1). Characterize the dynamics of the lower atmosphere (below about 75km) of Venus, including: retrograde zonal super-rotation, meridional circulation, radiative balances, mountain waves, and transfer of angular momentum.
		UD. Upper Dynamics (1). In the upper atmosphere and thermosphere of Venus, characterize global dynamics and interactions between space weather and the ionosphere and magnetosphere.
		MP. Mesoscale Processes (2). Determine the role of mesoscale dynamics in redistributing energy and momentum throughout the atmosphere of Venus.
	B. What processes determine the baseline and variations in Venus atmospheric composition and global and local radiative balance?	RB. Radiative Balance (1). Characterize atmospheric radiative balance and how radiative transport drives atmospheric dynamics on Venus.
		IN. Interactions (1). Characterize the nature of the physical, chemical, and possible biological interactions among the constituents of the Venus atmosphere.
		AE. Aerosols (2). Determine the physical characteristics and chemical compositions of aerosols in Venus atmosphere as they vary with elevation, including discrimination of aerosol types/components.
		UA. Unknown Absorber (2). Characterize the unknown short-wavelength absorber in the upper atmosphere of Venus and its influence on local and global processes.
		OG. Outgassing (3). Determine the products of volcanic outgassing on Venus and their effects on atmospheric composition.
	III. Understand the geologic history preserved on the surface of Venus and the present-day couplings between the surface and atmosphere.	A. What geologic processes have shaped the surface of Venus?
GC. Geochemistry (1). Determine elemental chemistry, mineralogy, and rock types at localities representative of global geologic units on Venus.		
GA. Geologic Activity (1). Characterize current volcanic, tectonic, and sedimentary activity that modifies geologic units and impact craters and ejecta on Venus.		
CR. Crust (2). Determine the structure of the crust of Venus in three dimensions and thickness across the surface.		
B. How do the atmosphere and surface of Venus interact?		LW. Local Weathering (1). Evaluate the mineralogy, oxidation state, and changes in chemistry of surface-weathered rock exteriors at localities representative of global geologic units on Venus.
		GW. Global Weathering (2). Determine the causes and spatial extents of global weathering regimes on Venus.
		CI. Chemical Interactions (3). Characterize atmospheric composition and chemical gradients from the surface to the cloud base both at key locations and globally.

3. Descriptions of the Goals, Objectives, and Investigations

3.1. Goal I: Understand Venus' early evolution and potential habitability to constrain the evolution of Venus-size (exo)planets.

Venus and Earth possibly both hosted liquid water oceans for billions of years—or alternatively these celestial cousins may have trod distinct evolutionary paths from the birth of the Solar System. Precisely because it may have begun so like Earth, yet evolved to be so different, Venus is the planet most likely to yield new insights into the conditions that determine whether a Venus-sized exoplanet can sustain long-lived habitability

Objective I.A. / Did Venus have temperate surface conditions and liquid water at early times?

The amount of water that Venus received during and after its accretion remains unknown. Standard models imply that Venus and Earth received similar amounts of water from comets and bodies that formed in the vicinity of Jupiter. Liquid water may have never condensed on Venus if its atmosphere was always hot and dense. Temperate surface conditions represent only one possible evolutionary path for Venus. However, this scenario is uniquely compelling given the implications for the habitability of ancient Venus and Venus-sized exoplanets at present day.

Investigation I.A.HO. Hydrous Origins (1)

This Investigation is considered Essential because of the importance of water to the geologic evolution and potential habitability of Venus. Although liquid water is not now present, water is required to form large amounts of some types of rock. A prime example would be abundant granitic rock, as has been suggested for some tesserae. On Earth, large continents only formed because water was relatively abundant in the crust and mantle where magmas were being generated. Similarly, some sedimentary rocks cannot form without liquid water, such as those rich in minerals such as sulfate and halide (evaporite), silica (hot springs or hardpans), or carbonates. Even deposits of clastic sediments can preserve physical signatures of transport by liquid water (e.g., delta deposits observed from orbit on Mars).

Remote sensing and in situ analyses may reveal signatures of hydrous origins. Many granitic rock types have much lower VNIR emissivity than basaltic rocks, and these different signatures can be accessed through several spectral 'windows' (near 1 μm) in Venus' thick atmosphere even if substantial weathering has occurred. Similarly, low emissivity may reveal sediments rich in evaporites, silica, or carbonates. Emissivity could be measured from orbit, from aerial platform, or from the surface. The physical characteristics of clastic sedimentary systems may be discernable from orbital or aerial radar, given high-enough spatial resolution. Landers are required to provide the most detailed determinations of rock type and physical inter-relationships using high-resolution imagers and chemical analysis instruments (such as x-ray fluorescence, gamma ray spectrometry, or LIBS). Lander instruments may remove any surface rind from chemical weathering to measure the detailed mineralogic composition of the rock.

Investigation I.A.RE. Recycling (1)

Crustal recycling occurs when surface and near-surface material is introduced by subduction and/or delamination to the planetary interior, where it can participate in melt production and the chemical evolution of the lower crust and mantle. This Investigation is Essential because identifying widespread ongoing or ancient crustal recycling on Venus would have profound implications for our understanding of the thermal, chemical, geological, and

atmospheric evolution of the planet, and of terrestrial planets in general. Localized plume-induced subduction has been proposed to operate on Venus, and there is evidence of substantial lateral mobility of some parts of the crust. Crustal recycling is predicted to result in lavas with distinctive geochemical signatures, and numerous regions of tesserae on Venus have been hypothesized to correspond to continental-like material formed during an earlier era of crustal recycling. However, currently available radar image and topographic data are insufficient to determine whether these processes operate more widely and/or took place in the geological past.

No single type of observation by itself can definitively establish crustal recycling on Venus. Nonetheless, global radar images, topography, and gravity data at high resolution could help search for geomorphological evidence of crustal recycling—especially if augmented by imagery and topography at higher resolutions at areas of interest. In situ measurements by a landed platform of lava flows could test for chemical evidence for recycling (e.g., enrichment or depletion of incompatible elements such as Nb, Zr, rare-earth elements, etc.). Similarly, direct chemical analyses of tessera terrain would provide a test of that material corresponding to Earth-like continental crust (e.g., Si abundance). Finally, detections of seismic activity on Venus with orbital, aerial, or landed assets may help constrain models of recent or ongoing crustal recycling.

Investigation I.A.AL. Atmospheric Losses (2)

Atmospheric loss processes on Venus are the upper boundary condition for the evolution of its atmosphere so this Investigation is Important. The high D/H ratio (~100 times that of Earth's oceans) implies that Venus might once have held an ocean's worth of water but, somehow, lost most of it to space. Non-thermal escape driven by the solar wind is the most important process for atmospheric loss today because the escape velocity is too high for atmospheric escape to be driven by thermal or photochemical processes. Intense solar wind disturbances, such as those generated by co-rotating interaction regions (CIRs) and interplanetary coronal mass ejections (ICMEs), are known to increase atmospheric escape. Observations of Venus' ion outflow during solar disturbances show that the escape flux can increase by orders of magnitude, especially during ICME events. Additionally, changes in the interplanetary magnetic field (such as those associated with CIRs) lead to magnetic reconnection on the Venusian dayside that further drives atmospheric loss. Atmospheric loss is always present via ambipolar diffusion, a process which is much more efficient at Venus than any other terrestrial planet.

Loss processes have not been well characterized although they were addressed by PVO and VEX. Simultaneous observations of both the upstream solar wind and the Venusian thermosphere and ionosphere over the full range of local times and solar zenith angles—during solar minimum and maximum—would build a more complete picture of atmospheric erosion. Relevant instruments for this Investigation include but are not limited to electron spectrometers, ion mass spectrometers, neutral particle detectors, UV and visible spectrographs and imagers, SEP detectors, Langmuir probes, and magnetometers. These measurements are best conducted with spacecraft (e.g., multiple SmallSats) that image dayglow and aura and sample plasma conditions, particularly ion acceleration, from the exobase (~150 km) out to several Venus radii (e.g., at the Sun-Venus L1 Lagrange point), as well as the upstream solar wind environment.

Investigation I.A.MA. Magnetism (3)

Venus has no intrinsic magnetism today but may have hosted a dynamo in the past because its rotation is fast enough that the Coriolis force would have a large effect on flow if the core were convecting. Detecting any crustal remanent magnetism would imply that a dynamo

existed in the past and that the surface has remained cooler than the blocking temperatures of magnetic carriers. For example, simulations predict that a dynamo may have operated within the average surface age if the core of Venus was initially “Earth-like,” meaning hot and chemically homogeneous. Common magnetic minerals such as magnetite and hematite may retain thermoremanent magnetization for billions of years at Venus temperatures. Data from Pioneer Venus Orbiter and Venus Express only rule out crustal magnetization that is both strong and has typical coherence wavelengths >150 km northwards of 50° South latitude. Venera 4 measured magnetic field down to ~ 25 km in Eistla Regio and failed to detect any crustal remanence.

Future orbiters could still detect fields produced by strong, large-scale crustal magnetization southward of 50° South latitude. Magnetization that is relatively weak and coherent over >150 km and/or strong and coherent over smaller scales could exist anywhere on Venus except the Venera 4 landing site. Magnetometer measurements at low altitudes such as from an aerial platform are vital because magnetic field power decreases rapidly (e.g., as distance cubed) at altitudes above the coherence wavelength of the source magnetization. While a non-detection would admit multiple scenarios (e.g., no dynamo and/or a hotter surface in the past), this Investigation is Targeted because any detection of crustal remanent magnetization would provide a unique constraint on recent climate history and atmospheric loss processes.

Objective I.B. | How does Venus elucidate possible pathways for planetary evolution in general?

Only two Venus-sized planets made of rock and metal exist in our Solar System, but myriad examples of Venus-sized exoplanets await discovery with new telescopes. No general model for the long-term evolution of terrestrial exoplanets can rest on a foundation of fundamental ignorance about Earth and Venus. While Objective I.A. focused on a uniquely compelling evolutionary scenario, the following Investigations consider many possibilities.

Investigation I.B.IS. Isotopes (1)

This Investigation is considered Essential because the isotopic composition of Venus’ atmosphere should preserve significant clues to the accretion, differentiation, and early evolution of Venus. The deuterium/hydrogen ratio has substantial uncertainties. Other atmospheric constituents such as N, C, Cl, and the heavy noble gases constrain the abundances, sources, and compositions of volatiles in Venus’ early atmospheres. The isotopic compositions of these gases will help define the sources of Venus’ volatiles (e.g., comets versus asteroids) and whether they were affected by atmosphere loss processes, the extent of outgassing from the solid planet, and (more speculatively) active biology. Xenon is critical because the terrestrial planets appear to have tapped distinct sources. The same processes that produced the high D/H ratio may have depleted Xe from the atmosphere. Radioactive decay of ^{40}K in the mantle produces ^{40}Ar , and thus measurements of atmospheric ^{40}Ar constrain the integrated amount of volcanic outgassing from the interior. Finally, measurements of oxygen isotopes could reveal whether Venus and Earth likely formed from the same reservoir of material. Similar isotopic ratios for Earth and Venus in comparison to Mars would relax a key constraint on models of the origin of the Moon.

Mass spectrometer measurements of atmospheric constituents are required to fulfill this Investigation. Analyses of material from depths where the atmosphere is well-mixed (i.e., deeper than the homopause) would be most useful. Any platform that enters the atmosphere—*aerial platforms or atmospheric skimmers, or probes and landers with heritage from Pioneer Venus or Venera/VeGa, respectively*—could potentially deploy a useful mass spectrometer.

Investigation I.B.LI. Lithosphere (1)

Venus shows no evidence for a global regime of plate tectonics similar to that observed on Earth, implying that the lithosphere does not sustain localized deformation over long spatial and temporal scales. The current dynamical regime of Venus—whether it involves stagnant-lid convection, heat pipes, episodic overturns, or another mode of mantle convection—remains poorly understood. Determining key lithospheric parameters is considered an Essential Investigation because the lithosphere provides the upper boundary condition for mantle convection and links interior activity with surface observables.

High-resolution radar imagery and altimetry with global coverage would permit quantitative analyses of stress state, tectonic faulting, and volcanism—which constrain lithospheric rheology. Electromagnetic sounding can determine whether the water content of the lithosphere is more or less than a few hundred parts per million. These datasets would constrain models of processes that support topography. Elastic thicknesses would be retrieved by modeling admittance and/or flexural bending seen in topography. Mechanical thicknesses would then be derived and thus used to estimate the thermal gradient in the lithosphere that prevailed when elastic flexure first occurred. Existing estimates of elastic thickness from gravity data often have large uncertainties due to the limited accuracy and resolution of the present gravity field. Values from modeling topographic bending are limited to small subset of volcanoes and coronae where flexure can be observed. Finally, seismology conducted from orbital, aerial, and/or landed platforms may determine the physical structure and deformation processes of the lithosphere.

Investigation I.B.HF. Heat Flow (2)

Models of mantle convection often predict that the total heat flow through the lithosphere to the surface of Venus is roughly half of the measured value for Earth at present day. However, these models remain unvalidated by direct observations. Measuring the lithospheric heat flow (e.g., thermal gradient times conductivity) would help evaluate models. This Investigation is Important as a benchmark for models of geodynamic evolution in synergy with other constraints.

In-situ measurements would provide the highest-quality heat flow measurements if the seasonal/daily temperature variation is quantified through coupled studies of surface brightness temperature. Selecting targeted locations that enable meaningful comparisons with models of mantle convection requires understanding the geologic history and relative ages of surface units. Electromagnetic sounding conducted from an aerial platform could determine lithospheric thermal gradients over larger areas. Lightning-caused Schumann resonances are capable of penetrating to depths of ~50–100 km if the lithospheric water content is less than a few hundred parts per million. Retrieving thermal gradients in shallow regions is potentially tractable even if the lithosphere is relatively wet. Any measurement of present-day heat flow and thermal structure would complement constraints on past heat flow from elastic thickness measurements.

Investigation I.B.CO. Core (2)

This Investigation is Important because the size and physical state of the core is a key constraint on models of the thermal evolution of Venus. Energy from giant impacts (kinetic) or rapid accretion (gravitational) is expected to completely melt the core as an initial condition. The relative sizes of the silicate mantle and the metallic core constrain their compositions and the thermodynamic conditions of accretion through the abundance of light elements (e.g., silicon and oxygen) in the core. Two basic measurements of Venus are missing: total moment of inertia and

radius of the core. Existing measurements of the tidal Love number also are arguably too imprecise to distinguish between a partially liquid core and one that has finished solidifying.

Orbiters with modern radio tracking would provide improved measurements of the tidal Love number with the required precision. The moment of inertia of Venus may be constrained even without spacecraft missions. Simultaneous observations of radar echoes with the Goldstone Solar System Radar and the Green Bank Telescope allow measurement of the instantaneous spin orientation of Venus. Tracking the instantaneous spin orientation over years and decades enables a measurement of the moment of inertia. As planned for the NASA InSight mission on Mars, a single station with radio science may measure the radius of the core given a sufficient lifetime.

3.2. Goal II: Understand atmospheric dynamics and composition on Venus.

The atmosphere of Venus is a planet-sized heat engine. Energy deposition and the efficiency with which that energy is distributed throughout the planet are key constraints on potential habitability. For Earth, a fleet of in situ and orbital platforms builds the complete, four-dimensional picture of atmospheric evolution. These Investigations divide the atmosphere of Venus into regional areas, but these areas ultimately remain coupled in a planetary system.

Objective II.A. / What processes drive the global atmospheric dynamics of Venus?

Many fundamental atmospheric characteristics of Venus are poorly understood, including the global cloud cover and retrograde zonal super-rotation (RZS). Winds on Venus flow primarily from east to west at almost all altitudes below about 85 km. Wind speeds reach a peak in magnitude just above the cloud tops. Above this ~75 km altitude, winds transition to a subsolar to antisolar (SSAS) flow, before transitioning back to RZS in the upper thermosphere.

Investigation II.A.DD. Deep Dynamics (1)

The existence of the super-rotation of Venus' atmosphere has been known from cloud top observations since the early 20th Century. Full understanding of the RZS structure in the Venus atmosphere and mechanisms for its maintenance remains elusive. Variability in the form of zonal jets has been inferred from Akatsuki observations. Global-scale waves observed by Akatsuki's Longwave Infrared camera are tied to the crests of continent-sized land masses and recur regularly at similar times-of-day. Similar waves may have been seen by the VeGa balloons near the dawn terminator while flying over Aphrodite Terra. These orographic waves demonstrate the importance of surface-atmosphere interactions for the dynamics of Venus and its atmosphere; generation and dissipation of the orographic waves has been inferred to produce measurable changes in the rotation rate of the solid planet, and to affect solar-atmosphere interactions even in the deep atmosphere. Finally, understanding RZS will be a milestone advance for atmospheric sciences in general, and provide tests of exotic behavior in models of exoplanetary atmospheres.

The deep atmosphere of Venus is defined here as the portion of the atmosphere lying beneath the cloud tops around ~75 km altitude. This Investigation is Essential because a critical transition in atmospheric dynamics occurs here according to measurements of wind speed and temperature. Above ~75 km altitude, the primary means of data acquisition must likely come from an orbital platform; while below this altitude, capability of measurement is expected to be possible from both in situ and remote means. Simultaneous observations of radar echoes with the Goldstone Solar System Radar and the Green Bank Telescope allow measurement of the instantaneous spin period of Venus, which was not achieved with either Magellan or Venus

Express. Tracking the instantaneous length-of-day provides a time history of variations in atmospheric angular momentum that can be used to constrain global circulation models.

Investigation II.A.UD. Upper Dynamics (1)

This Investigation is Essential to understand the RZS at planetary scales. The RZS flow begins to lose strength in the upper mesosphere with SSAS flow dominating near 100 km. However, RZS flow becomes prevalent again in the thermosphere for unknown reasons. Flow dynamics in the thermosphere may be constrained through observations of nightglow and auroral emission. Nightglow is atmospheric emission from photodissociated/ionized dayside atoms and molecules that are transported to the nightside and recombine to neutrals. Aurora refers to atmospheric emission due to particle precipitation across the nightside from the solar wind. One of the brightest Venusian nightglow features is the $1.27 \mu\text{m O}_2 (^1\Delta_g)$ emission which emits around 99 km near the antisolar point. VEX observations show that this nightglow feature traces the SSAS flow. However, simultaneous observations reveal that the NO UV bands (~115 km) are shifted three hours away from local midnight towards the dawn sector, indicating the recurrence of the RZS flow. PVO observations of auroral OI 130.4nm emission also exhibit an offset towards the dawn sector. Offsets are expected between the OI UV emission and the OI auroral emission that is present during solar storms, but no spatial mapping has been conducted.

Recent studies suggest that the ionosphere and neutral atmosphere are more intimately connected than previously believed. Measurements are needed of auroral and other excited gas emissions driven by solar processes as well as the solar wind in order to understand the connections between the solar wind and the Venusian atmosphere. Instruments needed to study the Venusian atmosphere and the solar wind via remote sensing and in-situ measurements include but are not limited to electron spectrometers, ion mass analyzers, ion and neutral mass spectrometers, energetic neutral atom detectors, UV and visible spectrographs and imagers, solar energetic particle (SEP) detectors, Langmuir probes, and magnetometers.

Investigation II.A.MP. Mesoscale Processes (2)

The previous two Investigations segregated global scale processes and observations according to vertical spatial domain (above and below an approximate altitude of 70 km, and the dominant processes in each). This Investigation is Important because smaller spatial scale processes can drive planetary atmosphere dynamics on either local or global scales. Adequate general circulation models of a planetary atmosphere must include reliable parameterizations of these “sub-grid-scale” processes, which mandates new observational constraints from Venus.

Critical mesoscale processes include the behavior and evolution of convective cells, horizontal and vertical wave propagation, and other mesoscale structures. These processes can be observed from orbit, as demonstrated by the discovery of numerous mesoscale features in both Venus Express and Akatsuki data. The rate and global distribution of lightning is a proxy for a certain rate of convective activity that is necessary to drive charge exchange. Direct in-situ measurement of the local dynamics of isolated convective structures and/or wave propagation would also contribute to this Investigation. However, spatio-temporal degeneracies would remain—as they have in the meteorological data from the VeGa balloons—in the absence of simultaneous measurements from orbiters and aerial platforms.

Objective II.B. / What processes determine the baseline and variations in Venus atmospheric composition and global and local radiative balance?

The atmosphere of Venus is a coupled chemical, radiative, and dynamical system. The composition and evolution of the atmospheric constituents are strongly regulated by chemical processes in the highly complicated, sulfur-based chemical networks. Yet, significant questions remain regarding the identities and/or the sources and sinks for many of these constituents.

Investigation II.B.RB. Radiative Balance (1)

This Investigation is Essential because gradients of the upwelling and downwelling radiative fluxes, in both incident solar and emitted planetary infrared, determine the heating/cooling rates that drive atmospheric dynamics. These gradients are determined by absorbers and scatterers of both shortwave and longwave radiation distributed throughout the atmosphere, and which are involved in a variety of physical and chemical interactions. Although the local radiative balance of Venus has been measured many times to reasonable precision (by the Pioneer Venus, Venera, and VeGa missions), a mismatch remains between the measured radiative and dynamical parameters of the Venus atmosphere and those produced by models. What is the magnitude of the influence of variability of the numerous radiatively active species, and cloud microphysics and opacity on the radiative balance of Venus? To what extent does this distribution of radiative sources and sinks drive the tropospheric dynamics?

Direct, in-situ measurement of spectrally resolved (or integrated) upwelling and downwelling radiances on both the nightside and dayside can support this Investigation. Probes, landers, and/or mobile aerial platforms could make relevant measurements. The utility of these measurements increases substantially with the number and lifetimes of platforms, and with global context. High spectral resolution, full spectrum measurements of emitted and reflected radiation would substantially contribute to this Investigation. Orbital assessments of radiative balance would benefit from improving the foundational models of atmospheric constituents.

Investigation II.B.IN. Interactions (1)

The constituents of the atmosphere of Venus are involved in a highly coupled system involving sulfur chemistry, organic chemistry, aerosol microphysics, and possibly even biological activity. This Investigation is Essential because agreement between the chemical models of the Venus atmosphere and the observed vertical profiles of multiple participating constituents is necessary for constraining models of radiative balance and atmospheric evolution. Examples of unresolved questions include the relative roles of OCS and SO₂ as sulfur donors to the sulfurohydrologic cycle of sulfuric acid generation, as well as the role of water in that process. The photochemistry and thermochemistry in the Venus atmosphere proceed in environments significantly dissimilar to Earth, which have not been fully explored in the lab. Finally, the chemistry of supercritical carbon dioxide in the deep atmosphere remains unconstrained, but recent research suggests that it could explain why the temperature profile measured by the VeGa probe is stable against convection in the lowermost atmosphere.

An orbiter capable of high spectral resolution measurements across a broad region of the spectrum and retrieving high precision vertical profiles of chemically relevant species could make substantial progress towards achieving this Investigation, especially when coupled with improved models of atmospheric chemistry in the Venus environment. In situ aerial platforms can make substantial progress on understanding the aqueous chemistry; much as has been done for Earth. Lightning has been mapped on the nightside by Pioneer Venus and at polar latitudes by Venus Express. Statistical assessments of the presence and distribution of lightning in the Venus atmosphere—beyond the limitations imposed by Akatsuki's orbit—would constrain the

effects of lightning discharges on atmospheric constituents. Finally, landers and descent probes capable of simultaneously measuring meteorological parameters and the mixing ratio of carbon dioxide (and other species) in the lowest ~10 km can study supercritical carbon dioxide.

Investigation II.B.AE. Aerosols (2)

This Investigation is Important because aerosols have a major impact on the Venus greenhouse effect as well as its remotely observable properties. Aerosols are an integral part of the atmospheric chemistry system both as an active constituent—for example, in the distribution and formation of sulfuric acid from sulfur-bearing compounds and water vapor—and as a passive constituent—for example, in the facilitation of heterogeneous chemistry on the surfaces and interiors of cloud droplets. While the primary constituent of the upper clouds has long been known to be spherical particles of highly concentrated sulfuric acid with typical radii of one micron, the exact nature of the Venusian aerosols is incompletely known. A submicron mode of particles is known to exist in both the upper clouds, as well as in and below the middle and lower clouds. However, the size distribution of this smallest mode of particles at all altitudes remains under-constrained by the available data. Furthermore, in the upper clouds, the composition has been assumed to be sulfuric acid, but this has never been definitively shown. In and below the lower and middle clouds, the composition of this submicron aerosol mode remains similarly unknown. Finally, the largest mode of particles, the “Mode 3” particles reported by the LCPS probes and VeGa landers in the middle and lower clouds, remains a controversial topic. The night side near infrared inhomogeneities are attributed largely to variations in the Mode 3 population, yet their existence remains unconfirmed and their composition unknown. Mode 3 particles may have a crystalline component, or merely represent the largest “Mode 2” particles.

In-situ nephelometer and mass spectroscopy of cloud aerosols would reduce uncertainties in aerosol size distributions and compositions. Observations at altitudes throughout the cloud column are key because different populations of aerosols occur at different altitudes.

Investigation II.B.UA. Unknown Absorber (2)

Short-wavelength visible and near-ultraviolet light is unaccountably absorbed in Venus’ upper atmosphere. The effects of this unknown absorber are strongest in the near-ultraviolet, but are apparent well into the wavelengths of visible light, and may have contributions elsewhere in the spectrum. The unknown absorber varies in strength over space and over a wide range of timescales. This Investigation is Important because the unknown absorber is responsible for at least half of the deposition of solar insolation into the atmosphere.

Numerous candidates have been proposed, including sulfur allotropes (S_x), iron chlorides ($FeCl_3$), and OSSO and its isomers. For example, OSSO provides a good spectral match to the UV absorption, and the high concentration of SO in the atmosphere implies that OSSO dimers should be present. However, the unknown absorber might not be a single species because OSSO does not produce sufficient absorption near 400nm to match observations. Other candidates such as S_4 might be expected to provide the necessary opacity in this spectral region. A biological origin for the unknown absorber in the “habitable zone” of the atmosphere has been suggested on the basis of comparison to the spectral properties of terrestrial acidophilic organisms.

Mass spectrometry of atmospheric material in the region of the unknown absorber is the measurement most likely to accomplish this Investigation. The platform most appropriate to carry this mass spectrometer is likely a descent probe, or aerial platform. In addition, high-

resolution spectroscopy from orbit or from an aerial platform can contribute to the investigation. In general, synergistic measurements from multiple platforms are desired.

Investigation II.B.OG. Outgassing (3)

This Investigation is Targeted to determine whether volcanic outgassing may provide insights into the composition of crust and the structure and dynamics of the interior. Indirect observations hint that Venus is volcanically active today. Overall, the surface has a young average age of ~750 Myr and hosts myriad volcanic surface features. Transient and high concentrations of SO₂ in the atmosphere and thermal anomalies on the surface have also pointed to currently active volcanism. Massive H₂SO₄/H₂O clouds are most likely the products of volcanic outgassing of SO₂ in the past ~10–50 Myr. The VIRTIS instrument on Venus Express measured near-IR surface emissivity anomalies interpreted as a lack of surface weathering at fresh volcanic flows on and near several massive shield volcanoes on Venus.

Direct imaging of hot spot volcanism and volcanic lakes at near-IR wavelengths and SAR interferometry from orbit have been suggested as means to characterize the ongoing rate of volcanism on Venus. Direct monitoring of heat from volcanic activity remains a viable, low-risk method to detect the ‘smoking gun’ of active volcanism on Venus. Spectroscopic remote sensing of transient gases in a volcanic plume (SO₂, H₂O) has also been suggested as an indirect means of sensing volcanic activity and outgassing. Aerial platform measurements acquired from beneath the clouds could be used to confirm the nature of events detected from orbited in images with five orders of magnitude better spatial resolution. In situ chemical measurements, including light isotope abundances, could help constrain the composition of outgassed materials.

3.3. Goal III: Understand the geologic history preserved on the surface of Venus and the present-day couplings between the surface and atmosphere.

Unveiling the past requires understanding the present. Although the NASA Magellan mission provided the first global maps of Venus, many first-order questions regarding their interpretation and implications await answers, which motivates collecting higher-resolution imagery, topography, and many other datasets that are available for other terrestrial planets.

Objective III.A. / What geologic processes have shaped the surface of Venus?

Since the Magellan mission, models for the geologic history of the surface have been divided between catastrophic and uniformitarian end-member scenarios. Moving beyond this controversy requires answering many basic questions about the present-day surface, including its stratigraphic history, composition, and potential for ongoing geologic activity.

Investigation III.A.GH. Geologic History (1)

Developing a stratigraphic history for the sequence of geological events is crucial for Venus to provide a framework for understanding the processes that shaped the coupled evolution of the surface, interior, and atmosphere. Volcanism and tectonism are ultimately driven by processes in the interior of Venus, and volcanism also contributes to development of the atmosphere. Similarly, the history of tectonic activity constrains the style and temporal evolution of convection in the mantle. In addition, a stratigraphic history for Venus would facilitate comparisons with other terrestrial planets. This Investigation is therefore considered Essential.

The key data sets for this Investigation are high-resolution radar imagery and topography. Magellan provided near global radar imaging at ~120–300 m/pixel and altimetry at ~8–25

km/pixel. Ideally, a follow-up mission will provide order-of-magnitude improvement in resolution in both radar images and altimetry. Global imaging coverage is desired to fully picture Venus' geologic evolution. A targeted survey could focus on mapping at high spatial resolution highland regions with obvious tectonic or volcanic features plus a representative fraction of the regional plains. The signal-to-noise ratio should be sufficiently large to detect variability that may be present in the radar dark plains. Infrared imaging of selected regions at relatively high spatial resolution from a variable altitude aerial platform, deployed below the cloud deck, would be complementary to the greater spatial coverage possible from an orbiter.

Investigation III.A.GC. Geochemistry (1)

This Investigation is Essential because the surface chemistry and mineralogy of Venus remains poorly characterized. Chemical analyses provided by the Venera and VeGa missions, although engineering and scientific triumphs, do not permit detailed interpretations, such as are routine for terrestrial analyses and MER APXS rover analyses from Mars. In particular, their XRF analyses of major elements did not return abundances of Na, and their data on Mg and Al are little more than detections at the $\sim 2\sigma$ level. Their analyses for K, U, and Th (by gamma rays) are imprecise, except for one (Venera 8) with extremely high K contents ($\sim 4\%$ K_2O) and one (Venera 9) with a non-chondritic U/Th abundance ratio. The landers did not return data on other critical trace and minor elements, like Cr and Ni. In addition, the Venera and VeGa landers sampled only materials from the Venus lowlands—they did not target sites in any of the highland areas: shield volcanos, tesserae, nor the unique plateau construct of Ishtar Terra. Given all these ambiguous results, rock types that may indicate igneous provenance cannot be identified. Similarly, no information is currently available to identify Venus mineralogy.

Investigation III.A.GA. Geologic Activity (1)

This Investigation is Essential because the relatively young surface age signals that Venus is an active planet. Key data sets are radar imaging and topography as well as seismic measurements. Comparison of radar imagery and altimetry from a future orbital mission with archival data from Magellan could detect surface changes over a period of several decades. Differential InSAR altimetry from a future orbiter could detect small topographic changes (≤ 10 cm vertically) due to active tectonism or volcanism but would be limited to detecting changes that occur over the timescale of a future mission. Seismic measurements via a long-lived lander of seismicity induced by active tectonism or volcanism would also be valuable. Measurements by a single lander would be sufficient to detect such activity, but measurements by a network would enable more quantitative analysis of the activity. Because the rate of such activity is not known, this approach is enhanced by increasing the duration of the seismic measurements. As demonstrated by Venus Express, fresh flows (i.e., with little chemical weathering) are observable with NIR spectroscopy from orbit. Experiments suggest these flows may be only years old.

Several types of supporting measurements also are possible. The coupling of surface motion into the thick Venus atmosphere can produce propagating waves in the upper atmosphere that are detectable in high temporal resolution infrared images from orbit or by infrasound measurements on an aerial platform. This is considered a supporting measurement because such measurements do not contain as much detail as the original seismic waveform. Volcanic flows temporarily raise the surface temperature, which could be measured by infrared or microwave radiometry. Observing the thermal signature requires flows with high flux rates, which prevent the flow from crusting over. Finally, outgassing associated with large explosive volcanic

eruptions may temporarily create a disequilibrium in the atmospheric composition which could be measured by orbiters, atmospheric entry probes or sondes, and/or aerial platforms.

Investigation III.A.CR. Crust (2)

This Investigation is Important because the crust of Venus has at least partially recorded the last billion years or so of tectonic and volcanic activity on the planet. Crustal thickness can constrain the total amount of magmatism, and variations related to location of more ancient materials like tessera or more recent units like rift zones can help quantify activity outside of the resurfacing event. Information about the structure of the crust, including the thickness of plain units and the penetration of faults at depth, are also crucial for reconstructing the history of geological activity on Venus and how it may have changed over time. The timing and volume of volcanic flows and their interactions with impact craters would constrain whether volatiles were released gradually or catastrophically from the interior.

Global radar images, topography, and gravity data collected by orbiters and/or aerial platforms at high precision and resolution would constrain the thickness and density structure of crustal units such as regional plains and volcanic flows. Improving the spatial resolution of global geological maps by one or more orders-of-magnitude would enable the delineation of individual lava flows, mapping individual fault blocks, and characterizing geologic contacts between volcanic and structural units, fundamentally transforming our understanding of volcanic and tectonic processes on Venus. Transmitting and receiving circular polarizations, as done by Arecibo and other planetary radars, would help quantify the thickness and grainsize of surface materials (e.g., volcanic deposits versus ejecta and/or regolith) Other geophysical techniques such as ground-penetrating radar (from orbital or aerial platforms) and seismology (from a surface instrument or detected in the atmosphere, which is strongly coupled to the ground) would provide strong constraints on the thickness and distribution of near-surface units on Venus.

Objective III.B. / How do the atmosphere and surface of Venus interact?

Temperatures of ~470°C and pressures ~90 bars near the surface ensure geologically rapid chemical reactions. Available data suggest that the deep atmosphere composition is not consistent with chemical equilibrium. However, significant uncertainties remain in the reactions that occur at the atmosphere-surface interface, the redox state of the atmosphere-surface boundary, and the concentrations and spatial variations of important trace gases near the surface.

Investigation III.B.LW. Local Weathering (1)

The history of gas/fluid interactions between Venus' hot, dense CO₂-rich atmosphere and its surface materials is recorded in the minerals that have experienced such alteration. Laboratory and phase equilibria studies predict oxidation of primary igneous minerals to oxide phases such as magnetite and hematite. This investigation would search for the presence of dehydrated minerals such as anhydrite and well as the possible presence of alteration phases from basaltic minerals. Such investigations could involve in situ instruments for mineralogy, including visible-mid-IR and Raman spectroscopies. Penetrating beneath surface alteration could provide valuable information on the depth of alteration as well as the underlying mineralogy.

Investigation III.B.GW. Global Weathering (2)

Among the most striking findings of the Magellan mission was the discovery of great differences in radar backscatter brightness with elevation: the highlands are significantly brighter

than the lowlands. However, approaches to identifying candidate substances responsible for this dichotomy depend on assumed surface geochemistry and oxidation state, both of which are poorly known. Possible investigations to resolve these questions are two-fold. Orbital spectroscopy utilizing the windows in the ~1-micron region make it possible to distinguish key rock types (e.g., basalt vs. granite) and can differentiate among phase changes with elevation, such as magnetite hematite, and pyrite. Surface mineralogy could also be measured using in situ visible-mid-IR spectroscopy, X-ray diffraction, and Raman Spectroscopy, while Mossbauer spectroscopy could measure oxidation state.

Investigation III.B.CI. Chemical Interactions (3)

This Investigation is Targeted to determine the abundances of crucial species (e.g. CO, OCS, SO₂) in the lowest atmospheric scale height where surface-atmosphere interactions will occur. Inferences about their near-surface concentration have previously been made through extrapolation of their observed higher elevation concentrations and through model predictions. Because the average Venus atmosphere is oxidized compared to basaltic rock, surface chemistry should produce reduced gas species, like CO from CO₂, and SO₂ or S₂ from SO₃. Venera landing sites indicate some rocks at the Venusian surface may be enriched in S relative to Earth basalts, suggesting processes of basaltic weathering. Second, SO₃ in the atmosphere may react with Ca-bearing silicates to form CaSO₄ (anhydrite) thus reducing the proportion of atmospheric sulfate. Atmospheric halogens could exchange with the surface, perhaps reducing the Cl/F ratio by formation of Cl-bearing phosphate phases. Third, atmosphere-surface chemistry varies with elevation in equatorial and northern highlands. Finally, if Venus' volcanic rocks include hydroxy-bearing igneous minerals (such as amphibole or biotite), then their decomposition should release hydrogen (with D/H values of the interior) to the atmosphere.

In-situ direct measurements of deep Venus atmosphere would provide clarity to questions of the concentrations and distributions of gases whose lowest scale height concentrations have only been inferred. This could be accomplished via descent probes with mass spectrometers designed to sample Venus' lower atmosphere as the probe descended and then for as long as they survive on the surface. Determining gradients on a regional scale would be enabled by orbital or aerial platforms carrying high-spectral-resolution spectrometers. Interpretation of these deep atmosphere spectra would be improved by better laboratory and/or theoretical estimates of the effects of pressure broadening on the specific line widths and strengths relevant to the Venus lower atmosphere. Additionally, experiments at the relevant temperature and pressure of the Venus surface environment could answer questions of which surface-atmosphere chemical reactions are plausible explanations for observed gas concentrations.

4. Conclusions

Despite being Earth's nearest neighbor, many fundamental questions about the origin and evolution of Venus await answers. Venus could have maintained a habitable environment with liquid water oceans for billions of years, and detectable signatures of this ancient epoch could still await detection today. Unveiling Venus is especially pressing given the rapid rate of ongoing discoveries of terrestrial exoplanets that may soon be amenable to atmospheric characterization. After a year-long process featuring a plethora of forums for community feedback, VEXAG has prepared this report centered on three Goals, six Objectives, and twenty-three Investigations that could drive a sustained program of Venus exploration. Dramatic scientific advancements are achievable within the next few decades if we can muster the collective will to explore Venus.

Appendix 1: Future Investigations

All Investigations in Table 1 were judged to have very high scientific merit along with feasibility in terms of technology and mission opportunities within the scope of this reports. However, this report neglected potential Investigations that may have very high scientific merit but relatively low feasibility. In particular, two Investigations were considered but judged to require resources substantially beyond the Flagship mission class and/or technology development outside the scope of these reports. We provide their descriptions to partially illustrate the types of science that would be become achievable after a program of Venus exploration has advanced:

Investigation III.A.AA. Absolute Ages

In the absolute sense, nothing is known about the surface age of rocks on Venus' surface. Although impact ages suggest it may be quite young, the possibility remains that some units might date from a time when Venus was habitable. Technology for in situ age dating is rapidly evolving, as evidenced by the success of the Sample Analyzer at Mars (SAM) instrument on Mars Science Laboratory. A long-term goal of the Venus Exploration Program is to obtain analogous in situ measurements of multiple locations on the surface. Current technology in development for this purpose includes SAM-like instruments and other solutions using high resolution LIBS. The latter method measures the emission spectra of molecules and molecular ions, enabling identification of specific isotopes within the plasma plume. Because sample preparation is not needed, LIBS provides a viable solution for Venus exploration.

Investigation I.B.DS. Deep Structure

Decades of study has revealed heterogeneous structure within Earth such as mantle plumes, laterally varying depths of seismic velocity discontinuities associated with mantle phase transitions, large low shear velocity provinces, and ultra-low velocity zones in the mantle. Seismology has also revealed hints of slow layers at the top and bottom of the liquid, outer core. This structure reflects thermal and/or compositional variations that constrain planetary accretion, differentiation, and ongoing processes. Considerable investment and technological development would be required to return Earth-quality seismic data from Venus. However, many signatures of important dynamical processes are likely buried in the deep interior. Excitingly, True Polar Wander (TPW) may occur quite rapidly on Venus relative to Earth and Mars because the equatorial bulge in the solid body is tiny and provides little obstacle to rotational realignment. Mantle convection (or, more detectable, large volcanic eruptions) could provide enough mass redistribution to provoke an episode of TPW.

Orbiter missions that conduct radar imaging with long temporal baseline could track the motion of surface features associated with TPW (e.g., at rates of ~1 m per year). Obtaining detailed constraints on plume structure, mantle seismic discontinuities, and chemical stratification in the lower mantle and core would require a global network of surface platforms.

Appendix 2: Linking the 2019 and 2016 VEXAG GOI Documents

The following table illustrates the connections between the investigations in this document and previous versions. Overall, items in the GOI have been reworded and reorganized but the overall scientific content remains mostly unchanged. Removing the relative prioritization of Objectives and Investigations is perhaps the most impactful difference between this document and previous versions. Because so many pressing questions about Venus await answers, it is accurate to describe multiple investigations as having the highest level of scientific priority.

Table A2.1. Investigations in the 2019 and 2016 VEXAG GOI	
Note that Objectives and Investigations were prioritized in the 2016 GOI. Investigations in the 2019 GOI are categorized but not prioritized within each category. For example, Investigations I.A.HO. and I.B.IS. have equal (highest) priority according to this 2019 GOI.	
Investigation in 2019 GOI	Related Investigation(s) in 2016 GOI
I.A.HO. Hydrous Origins (1)	III.A.2. III.A.3.
I.A.RE. Recycling (1)	II.A.3.
I.A.AL. Atmospheric Losses (2)	I.A.2.
I.A.MA. Magnetism (3)	II.A.3.
I.B.IS. Isotopes (1)	I.A.1. I.A.2. II.A.2. III.A.1. III.B.1. III.B.4.
I.B.LI. Lithosphere (1)	II.A.3.
I.B.HF. Heat Flow (2)	
I.B.CO. Core (2)	II.B.4.
II.A.DD. Deep Dynamics (1)	I.B.1. I.B.3. I.C.1.
II.A.UD. Upper Dynamics (1)	I.B.1. I.B.3.
II.A.MP. Mesoscale Processes (2)	I.B.1. I.B.3. I.C.1.
II.B.RB. Radiative Balance (1)	I.B.2.
II.B.IN. Interactions (1)	I.C.1.
II.B.AE. Aerosols (2)	I.C.1. I.C.2. I.C.3.

VEXAG Goals, Objectives, and Investigations for Venus Exploration (2019)

II.B.UA. Unknown Absorber (2)	I.C.2. I.C.4.
II.B.OG. Outgassing (3)	III.B.4.
III.A.GH. Geologic History (1)	II.A.1. II.B.6.
III.A.GC. Geochemistry (1)	II.B.1. II.B.2. II.B.5.
III.A.GA. Geologic Activity (2)	II.A.4.
III.A.CR. Crust (2)	II.B.3. II.B.6.
III.B.LW. Local Weathering (1)	III.B.2.
III.B.GW. Global Weathering (2)	III.B.2.
III.B.CI. Chemical Interactions (3)	III.B.3.

Appendix 3: Linking the 2019 VEXAG GOI and Roadmap

The VEXAG “Roadmap for Venus Exploration” describes a program of Venus exploration featuring twelve mission modalities. The following tables define each modality and indicate those that are potentially useful to each Investigation. Because the VEXAG GOI is not designed to prescribe particular missions, the omnibus table is only intended as a general guide.

Platform Type/Subtype Time Frame	Description of platform and primary scientific objectives
ORBITER	Supports investigations from orbital vantage points optimized for the scientific objectives
Surface/Interior Near Term	Single spacecraft in a circular, low altitude, near polar orbit optimized for most investigations of the surface and interior including those involving radar imaging and topography, infrared mapping, and gravity.
Atmosphere/Ionosphere Near Term	Single spacecraft in an eccentric, long-period orbit optimized for atmospheric remote sensing (e.g., nadir and limb viewing) and in situ sensors of the ionosphere and induced magnetosphere.
SmallSat or CubeSat Near Term	Single or multiple spacecraft focused on highly targeted investigations requiring tailored orbits. May also provide relay, navigation support, and synergistic science for surface and aerial platform(s).
ATMOSPHERIC ENTRY	Supports experiments during a traverse or descent in the Venus atmosphere
Skimmer Near term	Skims the atmosphere, sampling the Venus atmosphere at a very high altitude and emerging from the atmosphere for sample analysis and data relay.
Probe Near term	Enters the atmosphere and descends to the surface but not designed to operate after impact. Would investigate atmospheric structure and compositions along a single profile as well as near-surface imaging.
Sonde Mid term	Deploys from an aerial platform that is already at the operational altitude. Sonde relays data through the aerial platform as it descends. Advanced versions could target surface features.
SURFACE PLATFORM	Supports experiments on the surface of Venus in the high temperature high pressure environments
Short-Lived Near term	Classic (e.g., Venera) lander capable of surviving on the surface for several hours. Various instruments could investigate elemental and mineral compositions of nearby rocks, including variations with depth.
Long-Lived, Pathfinder Mid term	Designed to operate for at least one Venus solar day (≥ 117 Earth days) on the surface. Measurements include temperature, wind velocity, and chemistry of major species and possibly demonstration of a seismic sensor.
Long-Lived, Advanced Far term	Capable of both short duration (one Earth day) investigations of the surface and longer-term investigations of the atmosphere, heat flow and seismicity of the planet through at least two Venus solar days.
AERIAL PLATFORM	Supports extended duration experiments in and from the atmosphere including sonde deployment
Fixed Altitude – Mid Cloud Near Term	Floats at a nominal altitude of ~55 km in day and/or night at temperature near 20 °C. Carried around the planet in six days by the RZS and conducting investigations of the atmosphere and interior.
Variable Altitude – Mid Cloud Mid Term	Controls altitude in the range ~50–60 km enabling compositional and structural investigations of different regions within the clouds enhancing the range of investigations of the atmosphere and interior.
Variable Altitude – Cloud Base Far Term	Controls altitude in the range ~40–60 km using passive thermal control systems to enable use of conventional electronics. Sensors in exposed locations must tolerate temperatures up to 150 °C.

VEXAG Goals, Objectives, and Investigations for Venus Exploration (2019)

VEXAG GOI			Roadmap Mission Modalities												
Goal	Objective	Investigation	Orbiter	Orbiter	Orbiter	Atmospheric Entry			Surface Platform			Aerial Platform			
			Surface/Interior	Atmosphere	SmallSat	Skimmer	Probe	Sonde	Short-lived	Long-lived (Pathfinder)	Long-lived (Advanced)	Fixed Altitude	Variable Altitude	Variable+ Altitude	
			Near-term	Near-term	Near-term	Near-term	Near-term	Mid-term	Near-term	Mid-term	Far-term	Near-term	Mid-term	Far-term	
I. Early evolution and potential habitability	Did Venus have liquid water?	I.A.HO. (1)	Supporting	Vital	Vital						Supporting				
		I.A.RE. (1)	Supporting								Supporting				
		I.A.AL. (2)		Supporting	Supporting										
		I.A.MA. (3)		Supporting	Supporting					Supporting	Supporting	Supporting			
	How does Venus inform pathways for planets?	I.B.IS. (1)				Supporting	Supporting	Supporting	Supporting						
		I.B.LI. (1)	Supporting	Supporting						Supporting	Supporting	Supporting	Supporting	Supporting	
I.B.HF. (2)		Supporting								Supporting					
II. Atmospheric dynamics and composition	What drives global dynamics?	II.A.DD. (1)		Supporting	Supporting			Supporting	Supporting		Supporting	Supporting	Supporting	Supporting	
		II.A.UD. (1)		Supporting	Supporting	Supporting									
		II.A.MP. (2)		Supporting	Supporting			Supporting	Supporting			Supporting	Supporting	Supporting	
	What governs composition and radiative balance?	II.B.RB. (1)		Supporting	Supporting			Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	
		II.B.IN. (1)		Supporting	Supporting			Supporting	Supporting		Supporting	Supporting	Supporting	Supporting	
		II.B.AE. (2)		Supporting	Supporting			Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	
		II.B.OG. (3)	Supporting	Supporting	Supporting			Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	
III. Geologic history and processes	What geologic processes shape the surface?	III.A.GH. (1)	Supporting											Supporting	
		III.A.GC. (1)	Supporting	Supporting	Supporting			Supporting	Supporting	Supporting				Supporting	
		III.A.GA. (2)	Supporting	Supporting	Supporting			Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	
		III.A.CR. (2)	Supporting	Supporting	Supporting					Supporting	Supporting	Supporting	Supporting	Supporting	
	Atmosphere and surface interactions?	III.B.LW. (1)	Supporting	Supporting	Supporting			Supporting	Supporting	Supporting		Supporting	Supporting	Supporting	Supporting
		III.B.GW. (2)	Supporting	Supporting	Supporting					Supporting		Supporting	Supporting	Supporting	Supporting
III.B.CI. (3)		Supporting	Supporting	Supporting			Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	Supporting	

Color Code	Meaning
Supporting	Vital: Mission modality enables measurements that are vital (either alone or in combination) to completing the investigation.
Supporting	Supporting: Mission modality enables measurements that substantially contribute to completing the investigation.

Appendix 4: List of Acronyms

Acronym	Definition
APXS	Alpha Particle X-Ray Spectrometer
CIR	Co-rotating Interaction Regions
GOI	Goals, Objectives, Investigations
ICME	Interplanetary Coronal Mass Ejections
InSAR	Interferometric Synthetic-Aperture Radar
IR	InfraRed
JAXA	Japan Aerospace Exploration Agency
LCPS	Large probe Cloud Particle size Spectrometer
LIBS	Laser-Induced Breakdown Spectroscopy
MER	Mars Exploration Rover
NIR	Near InfraRed
PVO	Pioneer Venus Orbiter
RZS	Retrograde Zonal Super-rotation
SAM	Sample Analysis at Mars
SAR	Synthetic-Aperture Radar
SEP	Solar Energetic Particle
SSAS	SubSolar to AntiSolar
TPW	True Polar Wander
UV	UltraViolet
VeGa	Venera-Galley
VEX	Venus EXpress
VEXAG	Venus EXploration Analysis Group
VIRTIS	Visible and InfraRed Thermal Imaging Spectrometer
VNIR	Visible and Near InfraRed
XRF	X-Ray Fluorescence